

Discovery of VHE γ -ray emission from the BL Lac object B3 2247+381 with the MAGIC telescopes

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ABSTRACT

Aims. We study the non-thermal jet emission of the BL Lac object B3 2247+381 during a high optical state.

Methods. The MAGIC telescopes observed the source during 13 nights between September 30th and October 30th 2010, collecting a total of 14.2 hours of good quality very high energy (VHE) γ -ray data. Simultaneous multiwavelength data was obtained with X-ray observations by the *Swift* satellite and optical R-band observations at the KVA-telescope. We also use high energy γ -ray (HE, 0.1 GeV–100 GeV) data from the *Fermi* satellite.

Results. The BL Lac object B3 2247+381 ($z=0.119$) was detected, for the first time, at VHE γ -rays at a statistical significance of 5.6σ . A soft VHE spectrum with a photon index of -3.2 ± 0.6 was determined. No significant short term flux variations were found. We model the spectral energy distribution using a one-zone SSC-model, which can successfully describe our data.

Key words. BL Lac objects: individual(B3 2247+381)–galaxies:active–gamma rays

1. Introduction

The number of known extragalactic very high energy (VHE, $E > 100$ GeV) γ -ray emitters has increased from 6 to almost 50 (July 2011) during the past five years¹. Most of these objects are active galactic nuclei (AGN), especially blazars (flat spectrum radio quasars and BL Lac objects), which are known for their large variability across the electromagnetic spectrum from radio to VHE γ -rays. Eight of these new discoveries were made during high optical states of these sources, reported by the Tuorla blazar monitoring program, which have triggered MAGIC observations: Mrk 180 (Albert et al. 2006), 1ES 1011+496 (Albert et

al. 2007a), BL Lacertae (part of the data taken as ToO due to optical high state, Albert et al. 2007b), 3C 279 (Albert et al. 2008a, Aleksić et al. 2011a), MAGIC J0223+430 (Aliu et al. 2009), S5 0716+714 (Anderhub et al. 2009), B3 2247+381 (Mariotti et al. 2010, this paper) and most recently 1ES 1215+303 (Mariotti et al. 2011). This suggests that high optical states are an indication of a higher state also in VHE γ -ray band, at least in some sources. Additionally for PG 1553+113 a long term study (Aleksić, J., Alvarez, E.A., Antonelli, L.A., Antoranz, P. et al., ApJ, in prep.) suggests a correlation between the optical and VHE γ -ray flux, while for PKS 2155-304 it seems that, at least, in some cases the two wavebands are correlated (Foschini et al. 2007; Aharonian et al. 2009).

¹ <http://www.mpp.de/rwagner/sources/>

The object B3 2247+381 was selected from the sample presented in Nieppola et al. (2006) with a reported X-ray flux $F(> 1\text{keV}) > 2\text{ mJy}$. Donato et al. (2001) classify it as high-energy-peaked BL Lac object and the reported X-ray flux of the source is $F(> 1\text{keV}) = 0.6\text{ mJy}^2$. B3 2247+381 ($z = 0.119$, Falco et al. 1988) has been previously observed by the MAGIC telescope as part of the systematic search of VHE γ -rays from X-ray bright BL Lac objects (Aleksić et al. 2011b)³. MAGIC observed the source between August and September 2006, which resulted in an upper limit $F(> 140\text{GeV}) < 1.6 \cdot 10^{-11}\text{cm}^{-2}\text{s}^{-1}$. The source has been monitored in the R-band by the Tuorla blazar monitoring program ever since.

B3 2247+381 was also included in the list of potentially interesting TeV sources released to the Imaging Atmospheric Cherenkov Telescope experiments by the *Fermi*-LAT collaboration on 27th October 2009 (*Fermi*-LAT Collaboration 2009, priv. comm.). B3 2247+381 is listed in the first *Fermi*-LAT catalog of AGNs (Abdo et al. 2010) as 1FGL J2250.1+3825 with a very hard spectrum (spectral index of -1.6 ± 0.1). In the second *Fermi*-LAT catalog this object is listed as 2FGL J2250+3825 with a spectral index of -1.84 ± 0.11 (Ackerman, M., Ajello, M., Allafort, A., Antolini, E. et al., ApJ, in prep.). Neronov et al. (2010) found a hint of signal at 2.73σ in the *Fermi*-LAT data above 100 GeV over the period August 2008-April 2010.

In this paper we present the first detection of VHE γ -ray emission triggered by the optical high state of the source and the first optical light curve of B3 2247+381. We also present simultaneous X-ray data obtained by the *Swift* and High energy (HE) data obtained by *Fermi*-LAT satellites.

2. Observations and data analysis

2.1. MAGIC observations and data analysis

The VHE γ -ray observations were carried out with the MAGIC telescopes located on the Canary Island of La Palma (28.8° N , 17.8° W at 2200 m.a.s.l). The two 17 m telescopes use the atmospheric Cherenkov imaging technique and allow for measurements at a threshold as low as 50 GeV.

B3 2247+381 was observed with the MAGIC telescopes during 13 nights between September 30th and October 30th 2010 collecting a total of 21.2 hours of data, of which 5.3 hours were discarded based on the event rate. The effective time of this observation, correcting for the dead time of the trigger and readout systems is 14.2 hours. Part of the data was taken under moderate moonlight conditions.

All the data were taken in the false-source tracking (wobble) mode (Fomin et al. 1994), in which the telescope was alternated every 20 minutes between two sky positions at 0.4° offset from the source. The zenith angle was between 8° and 35° .

The data was analysed using the standard MAGIC analysis framework "MARS" as described in Moralejo et al. (2009) with additional adaptations incorporating the stereoscopic observations. The images were cleaned using timing information as described in Aliu et al. (2009) with absolute cleaning levels of 6 phe (so-called "core pixels") and 3 phe ("boundary pixels") for the first telescope and 9 phe and 4.5 phe respectively for the second telescope. The images were parameterized in each telescope separately following the prescription of Hillas (1985).

² Note, that Veron-Cetty & Veron (2006) list the source as a BL Lac candidate and that the classification has not been confirmed.

³ The stacked dataset of all X-ray selected blazars resulted in a significant γ -ray excess, which hints towards the fact that the sources were emitting at a very low level.

In order to reconstruct the shower arrival direction we used the random forest regression method (RF DISP method, Aleksić et al. 2010) which was extended using stereoscopic information such as the height of the shower maximum and the impact distance of the shower on the ground. We estimated the RF DISP for each telescope separately and obtained two possible solutions along the major axis of the shower image in each telescope, respectively. Finally, we searched for the combination of two solutions (one from each telescope) that have the shortest squared angular distance between them. If this squared angular distance is greater than 0.05 degree^2 the event is removed from further analysis. This improves the background rejection since hadron induced showers have a larger error on the reconstruction of the arrival direction. The final arrival direction is the mean of the two solutions (weighted by the number of pixels in each shower image).

For the gamma-hadron separation the random forest method is used (Albert et al. 2008b). In the stereoscopic analysis image parameters of both telescopes are used as well as the shower impact point and the shower height maximum. A detailed description of the stereoscopic MAGIC analysis can be found in (Aleksić J., Alvarez, E.A., Antonelli, L.A., Antoranz, P. et al., Astr. Particle Phys. in prep.).

2.2. Optical observations and data analysis

B3 2247+381 has been observed by the Tuorla group⁴, using the 35 cm KVA telescope, located on La Palma, since 2006 (see Aleksić et al. 2011a for a description of the telescope). Observations have been made in the R-band. The brightness of the object was measured based on stars in the same CCD-frames as B3 2247+381. These stars were calibrated by the Tuorla group during the observing seasons. R-band magnitudes were converted to fluxes, using: $F(\text{Jy}) = 3080 \cdot 10^{-m_R/2.5}$. The fluxes were corrected for galactic absorption by $R=0.398\text{ mag}$ (Schlegel et al. 1998). During the years 2006-2009 B3 2247+381 was a quite faint and steady source at $R \sim 1.8\text{ mJy}$, but during late summer 2010 it went to a high optical state, reaching an average flux level of 2.4 mJy . The source also stayed at this level throughout the observing season (see Figure 3). In late September an alert was given to MAGIC and it started observations on September 30th. An alert is issued when the optical flux has increased by 50% from the long term running average.

2.3. Swift observations and data analysis

The prime objective of the *Swift* Gamma-Ray Burst observatory, launched in November 2004 (Gehrels et al. 2004), was to detect and follow up on Gamma-Ray bursts, but it has turned into a multi-purpose observatory due to its fast slewing and response capacity and its multi-wavelength coverage. *Swift* is equipped with three telescopes, the Burst Alert Telescope (BAT; Barthelmy 2005), which covers the 15-150 keV range, the X-ray telescope (XRT; Burrows et al. 2005) covering the 0.3-10 keV energy band, and the UV/Optical Telescope (UVOT; Roming et al. 2005) covering the 1800-6000 Å wavelength range.

Swift Target of Opportunity observations were requested and from October, 5 to 16, 2010. *Swift*/XRT observed the source for $\sim 5\text{ ks}$ every night, in Photon counting (PC) mode. We also analysed *Swift* archival data from August 10th, 2009, February 18th, 2010 and April 18th, 2010 in order to compare the level of the X-ray emission to previous observations.

⁴ <http://users.utu.fi/kani/>

The data were processed with standard procedures using the FTOOLS task XRTPipeline (version 0.12.6) distributed by HEASARC within the HEASoft package (v.6.10). Events with grades 0-12 were selected for the PC data (see Burrows et al. 2005) and the response matrices available in the *Swift* CALDB (v.20100802) were used. For the spectral analysis, we extracted the PC source events in the 0.3-10 keV range within a circle with a radius of 20 pixels (~ 47 arcsec). The background was extracted from an off-source circular region of 40-pixels radius.

The spectra were extracted from the corresponding event files and binned using GRPPHA to ensure a minimum of 30 counts per bin in a manner so that the χ^2 statistic could reliably be used. Spectral analyses were performed using XSPEC version 12.6.0. We adopted both a simple power law model and a log-parabolic model as in Massaro et al. (2004) with an absorption hydrogen-equivalent column density fixed to the Galactic value in the direction of the source, namely $1.2 \times 10^{21} \text{ cm}^{-2}$. The two models provide similar results in terms of goodness of fit above ~ 0.7 keV. However below this energy the differences are in general negligible due to low statistics.

Swift/UVOT observed the source in the "filter of the day" mode, that is a different filter was used for different observations. This does not allow to compare the UV fluxes among different days. UVOT source counts were extracted from a circular region 5 arcsec-sized centered on the source position, while the background was extracted from a larger circular nearby sourcefree region. These data were processed with the *uvotmaghist* task of the HEASOFT package. The observed magnitudes have been corrected for Galactic extinction $E_{B-V} = 0.149$ mag (Schlegel et al. 1998).

2.4. Fermi data analysis

The *Fermi*-LAT is a pair conversion telescope designed to cover the energy band from 20 MeV to greater than 300 GeV (Atwood et al. 2009) which operates in survey mode, scanning the entire sky every three hours. The data used in this paper encompasses the time interval from August 5th, 2008 to April 7th, 2011 (MJD 54683 - 55658), and were analyzed with the *Fermi* Science Tools package version v9r23p0, which are available from the *Fermi* Science Support Center (FSSC). Only events with the highest probability of being photons, those in the diffuse class, located within 12° of B3 2247+304 were used in this analysis. In addition, a cut on the maximum zenith angle ($< 100^\circ$) was applied to reduce the contamination from the Earth-limb gamma-rays, which is produced by cosmic-rays interacting with the upper atmosphere. The background model used to extract the γ -ray signal includes a Galactic diffuse emission component and an isotropic component (including residual instrument background), modelled with the files `gll_iem_v02_P6_V11_DIFFUSE.fit` and `isotropic_iem_v02_P6_V11_DIFFUSE.txt`, which are publicly available⁵. The normalizations of the components comprising the total background model were allowed to vary freely during the spectral point fitting. The spectral fluxes were derived with the post-launch instrument response functions (IRF) `P6_V11_DIFFUSE`, and applying an unbinned maximum likelihood technique (Mattox et al. 1996) to events in the energy range spanning from 300 MeV to 300 GeV. All the sources from the 2FGL catalog (Abdo, A.A., Ackermann, M., Ajello, M., Allafort, A. et al., ApJS, in prep.) located within 7° radius of B3 2247+38 were

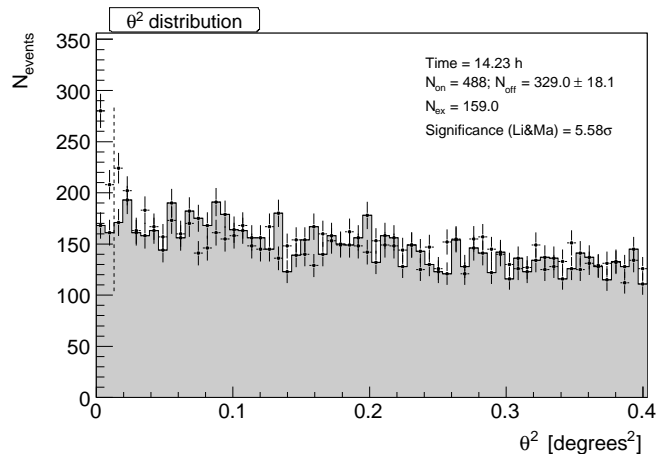


Fig. 1. Distribution of the squared angular distance (θ^2) for the on-source counts in the direction of B3 2247+381 (black points with errorbars) and the normalized off-source events (gray histogram).

included in the model of the region. The initial parameters in the XML file were set to those of the 2FGL catalog, leaving the normalization parameters free in the fitting procedure. The systematic uncertainty in the flux is estimated as 10% at 100 MeV, 5% at 560 MeV and 20% at 10 GeV and above⁶.

For the period of the MAGIC observations (30 days between September 30th and October 30th 2010), the source is not significantly resolved, and hence only 95% confidence level upper limits were obtained. In the light curve presented in Fig. 3, for each time bin, if the TS value for the source was $TS < 4$ or the number of predicted photons $N_{pred} < 3$, the values of the fluxes were replaced by the 95% confidence level upper limits. The 2-sigma upper limits were computed using the Bayesian method (Helene 1983), where the likelihood is integrated from 0 up to the flux that encompasses 95% of the posterior probability.

3. Results

In the MAGIC data the distribution of the square of the angle θ between the reconstructed direction of the events after cuts and the position of B3 2247+381 (RA: 22.83472° , DEC: 38.41028° , J2000 as in Ficarra et al. 1985) shows an excess of 159 ± 28 events above a threshold of 200 GeV. The cuts were previously optimized on a trial sample of Crab Nebula data. The threshold was calculated from Monte Carlo simulated data by finding the peak of the differential rate vs. energy distribution after cuts and spectral re-weighting. The expected background level with the same cuts is 329 events, calculated from the θ^2 distribution of the reconstructed direction of the events with respect to the anti-source position, located 180° from the real source position in the camera plane (Figure 1). The measured excess corresponds to a post-trial significance of 5.6σ calculated using eq. 17 from Li and Ma (1983).

The source position and extension, determined by a 2D Gaussian fit of the sky map produced with the cuts above, are consistent with a point-like source placed at the position of B3 2247+381 within 0.015° , well within the statistical uncertainty and the systematic pointing uncertainties of MAGIC.

The integral flux of the source above 200 GeV was estimated to be $(5.0 \pm 0.6_{stat} \pm 1.1_{sys}) \cdot 10^{-12} \text{ ph cm}^{-2}\text{s}^{-1}$. The differen-

⁵ <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

⁶ http://fermi.gsfc.nasa.gov/ssc/data/analysis/LAT_caveats.html

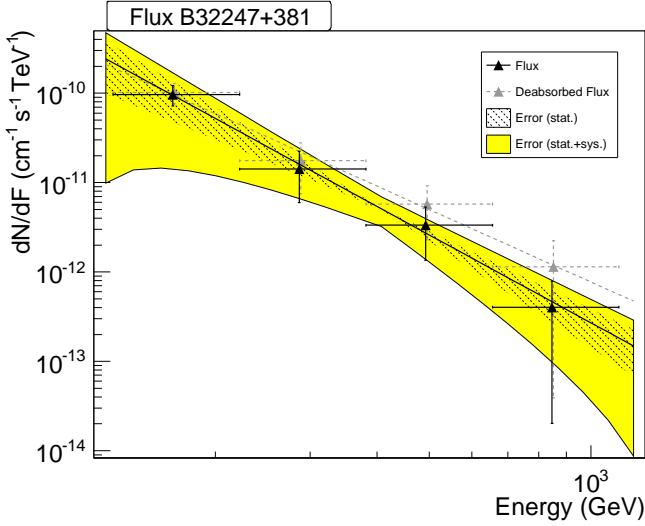


Fig. 2. The unfolded differential energy spectrum of B3 2247+381 observed by MAGIC. The black data points refer to the measured spectrum, while the grey dashed points have been corrected for the attenuation of the EBL. The thick black and dashed grey lines are power-law fits to the respective data points (fit results are given in the text). The dashed band corresponds to the statistical error of the fit to the measured spectrum, while the white band surrounding it is the sum of the statistical and systematic errors of the fit.

tial energy spectrum is well described by a simple power-law: $dN/dE = f_0 (E/300 \text{ GeV})^\gamma$. The photon index γ was found to be $-3.2 \pm 0.5_{\text{stat}} \pm 0.5_{\text{sys}}$, and a flux normalization f_0 at 300 GeV of $(1.4 \pm 0.3_{\text{stat}} \pm 0.2_{\text{sys}}) \cdot 10^{-11} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$. In order to correct the effects in the spectrum determination introduced by the limited energy resolution, different unfolding algorithms (Forward, Tikhonov, Schmelling 1&2, Bertero; all described in Albert et al. (2007c)) were used, and all agreed within errors. Taking into account the attenuation due to pair production with the extragalactic background light (EBL), the spectrum was found to be compatible with a power law with photon index $\gamma = -2.7 \pm 0.5_{\text{stat}} \pm 0.5_{\text{sys}}$ and flux at 300 GeV $f_0 = (2.0 \pm 0.3_{\text{stat}} \pm 0.3_{\text{sys}}) \cdot 10^{-11} \text{ ph cm}^{-2} \text{ s}^{-1} \text{ TeV}^{-1}$ (Figure 2). Two different EBL models were used (Domínguez et al. 2011 and Kneiske & Dole 2010), and they were found to be in good agreement with each other, well within the statistical uncertainties.

Long term light curves of B3 2247+381 in VHE γ -rays (MAGIC), HE γ -rays (*Fermi*-LAT), X-rays (*Swift*) and optical (Tuorla Observatory) are shown in Figure 3. Our detection in VHE γ -rays is compatible with the previous upper limit from 2006 and thus no variability can be established in this energy band. However, in X-rays and the optical band a clear increase of the flux in fall 2010 is evident, while the *Fermi*-LAT light curve is consistent with a constant flux. A fit with a constant to the eleven *Fermi* flux points where the source is significantly detected, gave a flux value of $(3.7 \pm 0.5) \cdot 10^{-9} \text{ ph cm}^{-2} \text{ s}^{-1}$, with a reduced χ^2 of 0.7 with ten degrees of freedom. The *Fermi*-LAT is not sensitive enough for detecting short term variations at this flux level. The temporal evolution of the VHE γ -ray, X-ray and optical emission from B3 2247+381 during September-October 2010 observation shows no strong variability on time scales of a night (Figure 4). In particular, the MAGIC light curve above 200 GeV is consistent with a non-variable source, having a reduced χ^2 of 0.6 with ten degrees of freedom. During one night

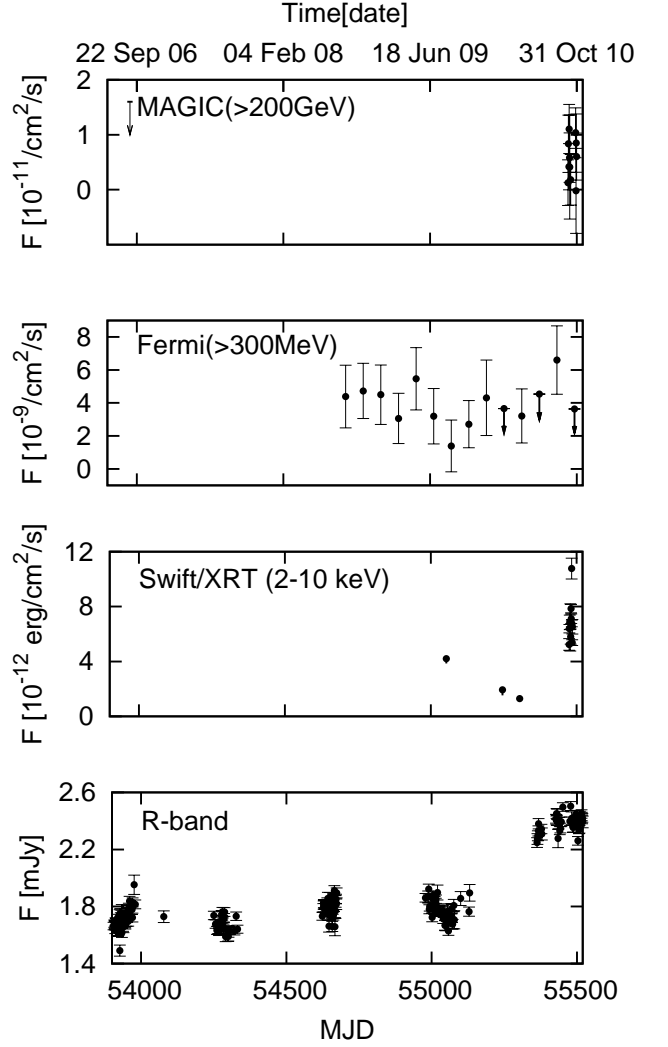


Fig. 3. Long term light curves of B3 2247+381 in VHE γ -rays, *Fermi*-LAT HE γ -rays (two months time intervals), *Swift* X-rays and optical KVA R-band.

the X-ray flux is significantly higher (almost factor 2) than the other X-ray points, but unfortunately we do not have simultaneous optical or MAGIC data for that night.

4. Discussion

In this paper we report the discovery of VHE γ -rays from B3 2247+381 by MAGIC. The MAGIC observations were triggered by an optical high state of the source, like several other discoveries. However, the observed VHE γ -ray flux is consistent with the upper limit from previous observations and we therefore cannot conclude if the source was in a higher VHE γ -ray state during the observations.

In Figure 5 we show the spectral energy distribution of the source during the MAGIC observations, together with simultaneous *Swift* and optical data, and other non-simultaneous data. The *Swift* observations show that the source was in a high state also in X-rays. In the *Fermi* energy range the source is very weak, which limits the capability of detecting statistically significant flux-variability on time scales of a few months. The synchrotron component of the SED is showing a significantly larger emission in the high state, while the inverse Compton compo-

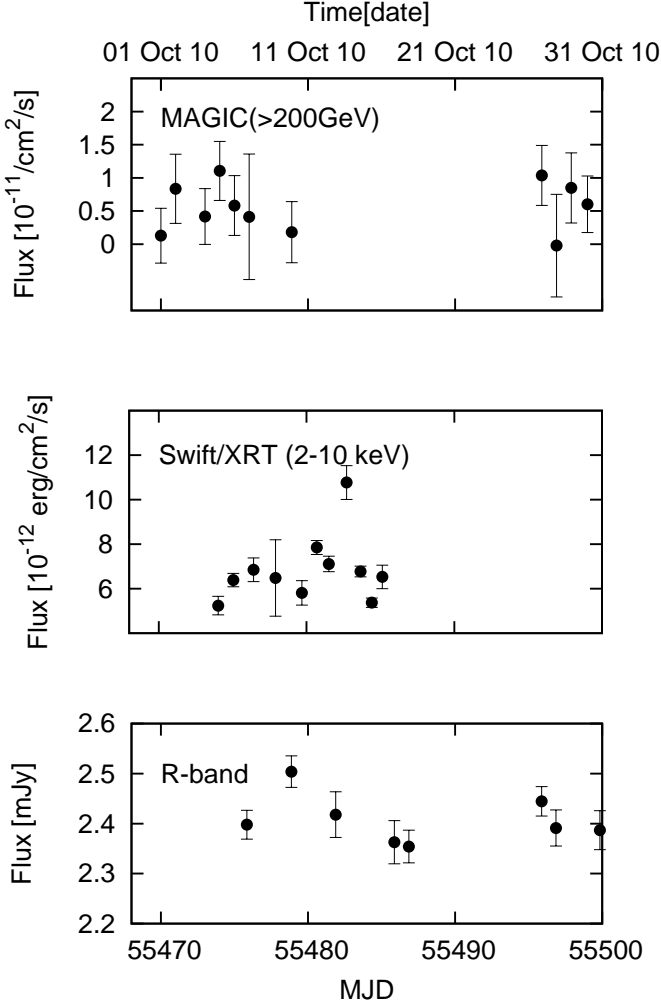


Fig. 4. Same as Fig. 3, but zoomed into the time interval of the MAGIC observations in September-October 2010.

ment is consistent with only minor changes. We must however note that the weak detection in the HE and VHE γ -ray band significantly limits the determination of the inverse Compton component.

We reproduce the SED with a one-zone synchrotron-self-Compton (SSC) model (see Tavecchio et al. 2001 for a description). In brief, the emission region is assumed to be spherical, with a radius R , filled with a tangled magnetic field of intensity B . The relativistic electrons follow a smoothed broken power-law energy distribution specified by the limits γ_{\min} , γ_{\max} and the break at γ_b as well as the slopes n_1 and n_2 before and after the break, respectively. Relativistic effects are taken into account by the Doppler factor δ . The used input model parameters are shown in Table 1.

To reproduce the change of the Compton and synchrotron luminosity ratio between the low and the high state we mainly act on the electron normalization, source radius and Doppler factor (with slight changes also to the other parameters). The steeper X-ray slope in the low state implies a larger value of n_2 .

The comparison with parameters derived for BL Lac objects (e.g. Tavecchio et al. 2010) reveals that the parameters used for B3 2247+381 are close to the typical values. As for other sources, the somewhat larger (lower) value of the Doppler factor

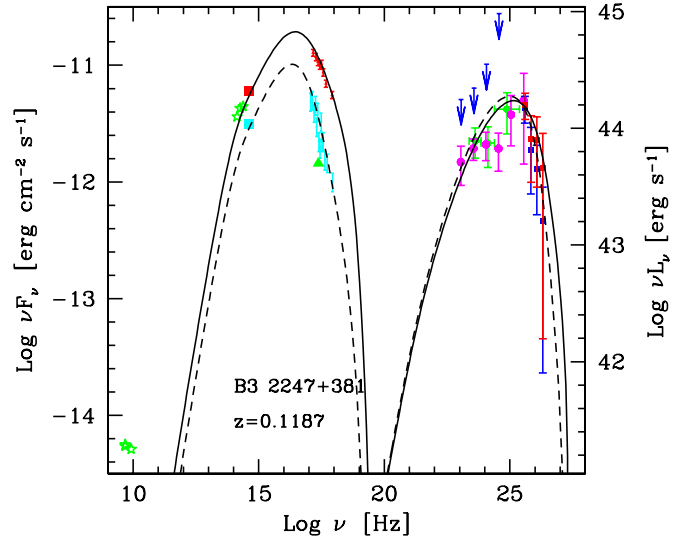


Fig. 5. Spectral energy distribution of B3 2247+381 (red: EBL corrected MAGIC spectral points). The green crosses are 1FGL *Fermi* data points (Abdo et al. 2010), while the pink points represent the *Fermi* analysis from this work (2.5 years of data). Blue arrows show the 95% confidence upper limits computed from *Fermi*-LAT data for the time interval of the MAGIC observation. Low (high) state *Swift* data were taken on April 18th 2010 (October 5-16, 2010). The host galaxy contribution has been subtracted from the KVA R-band data (red and light blue squares), following Nilsson et al. (2007). The data have been corrected for galactic absorption. Green and light blue points represent non-simultaneous low state data. The solid line is our SSC-model fit to the high state observations; the dotted line is a fit to the low state observations.

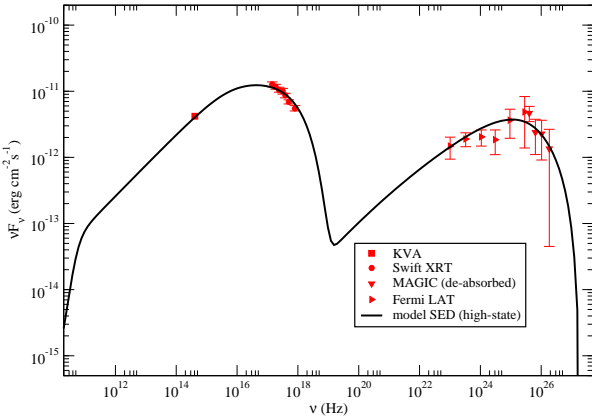
δ (the magnetic field intensity B) with respect to “standard” values is mainly due to the relatively large separation between the synchrotron and IC peaks.

Additionally, we model the SED, during the high state, with the one-zone SSC code from Weidinger et al. (2010) which is shown in Figure 6. In contrast to the model from Tavecchio et al. (2001) all parameters are basic physical parameters and the electron and photon spectrum and their breaks are derived self-consistently with a continuous injection of monochromatic electrons at the energy $\gamma_0 = 10^4$ and injection rate $K = 8.4 \cdot 10^4 \text{ cm}^{-3} \text{ s}^{-1}$. The spectrum is the resulting steady-state solution. The environment is defined by the magnetic field $B = 0.07 \text{ G}$, the acceleration efficiency $t_{\text{acc}}/t_{\text{esc}} = 1.09$ (i.e. the particle spectral index is $s = 2.09$ and the resulting $\gamma_{\max} = 4.8 \cdot 10^5$), and the blob radius $R = 1.3 \cdot 10^{16} \text{ cm}$. The break in the e^- spectrum of 1 at $\gamma_b = 2.9 \cdot 10^4$ arises self-consistently from IC and synchrotron cooling. The common parameters of both models agree very well.

The optical monitoring of candidate VHE γ -ray blazars has proven to be a successful tool for their discovery. However, for B3 2247+381 (like for Mrk 180 and 1ES 1011+496) we cannot establish a connection between the optical high state and the VHE γ -ray emission since the upper limit from previous observations (during a low optical state) is higher than the detected

Table 1. Input parameters for the high and low states of the SSC-model shown as solid and dashed lines in Fig. 5. For more explanations see text.

Flux State	γ_{\min}	γ_b	γ_{\max}	n_1	n_2	B G	K cm^{-3}	δ	R cm
High	$3 \cdot 10^3$	$7.1 \cdot 10^4$	$6 \cdot 10^5$	2.0	4.35	0.06	$2.5 \cdot 10^3$	35	$8 \cdot 10^{15}$
Low	$3 \cdot 10^3$	$6.8 \cdot 10^4$	$5 \cdot 10^5$	2.0	5.35	0.08	$1.15 \cdot 10^4$	30	$4 \cdot 10^{15}$

**Fig. 6.** The solid line shows the SED model using Weidinger et al. (2010). The data points are described in the inlay of the figure. The fit parameters can be found in the text.

VHE γ -ray flux during the discovery. Therefore further observations are still needed to study the connection between these two wavebands (see also Reinthal et al. 2011).

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References

Abdo, A.A. et al. 2010, *ApJ*, 715, 429
 Aharonian, F. et al. 2009, *ApJ*, 696, L150
 Albert, J. et al. 2006, *ApJL*, 648, L105
 Albert, J. et al. 2007a, *ApJL*, 667, L21
 Albert, J. et al. 2007b, *ApJL*, 666, L17
 Albert, J. et al. 2007c, *Nucl. Instr. Meth. A* 583, 494
 Albert, J. et al. 2008a, *Science* 320, 1752
 Albert, J. et al. 2008b, *Nucl. Instr. Meth. A* 588, 424
 Aleksić, J. et al. 2010, *ApJ*, 519, A32
 Aleksić, J. et al. 2011a, *A&A*, 530, A4

Aleksić, J. et al. 2011b, *ApJ*, 729, 115
 Aliu, E. et al. 2009, *ApJL*, 692, L29
 Anderhub, H. et al. 2009, *ApJL*, 704, L129
 Atwood, W.B. et al. 2009, *ApJ*, 697, 1071
 Barthelmy, S.D. 2005, *Space Science Reviews*, 120, 143
 Burrows, D.N. et al. 2005, *Space Science Reviews*, 120, 165
 Domínguez, A. et al. 2011, *MNRAS*, 410, 2556
 Donato, D., Ghisellini, G., Tagliaferri, G., Fossati, G. 2001, *A&A*, 375, 739
 Falco, E.E., Kochanek, C.S., and Munoz, J.A. 1988, *ApJ*, 494, 47
 Ficarra, A. et al. 1985, *A&AS*, 59, 255
 Fomin, V. P. et al. 1994, *ApJ*, 2, 137
 Foschini, L. et al. 2007, *ApJ*, 687, 81
 Gehrels, N. et al. 2004, *ApJ*, 611, 1005
 Helene, O. 1983, *Nuclear Instruments and Methods in Physics Research*, 212, 319
 Hillas, A. M. 1985, *Proc. of the 19th ICRC*, 3, 445
 Kneiske, T.M. & Dole, H. 2010, *A&A*, 515, 19
 Li, T.-P. & Ma, Y.-Q. 1983, *ApJ*, 272, 317
 Mariotti, M. et al. 2010, *Atel # 2910*
 Mariotti, M. et al. 2011, *Atel # 3100*
 Massaro, E. et al. 2004, *A&A*, 413, 489
 Mattox, J. R. et al. 1996, *ApJ*, 461, 396
 Moralejo, A. et al. 2009, in *Proc. 31st ICRC (Lodz)* (arXiv:0907.0943)
 Neronov, A., Semikoz D.V. & Vovk, I. 2010 (arXiv:1004.3767)
 Nieppola, E., Tornikoski, M., Valtaoja, E. 2006, *A&A*, 445, 441
 Nilsson, K. et al. 2007, *A&A*, 475, 199
 Reinthal, R. et al. 2011, *Proc. Beamed and Unbeamed Gamma-rays from Galaxies*, Open Access Journal of Physics: Conference Series (JCPs), in press
 Roming, P.W.A. et al. 2005, *Space Science Reviews*, 120, 95
 Schlegel, D. J., Finkbeiner, D. P., & Davis, M. 1998, *ApJ*, 500, 525
 Tavecchio, F. et al. 2001, *ApJ*, 554, 725
 Tavecchio, F., Ghisellini, G., Ghirlanda, G., Foschini, L., & Maraschi, L. 2010, *MNRAS*, 401, 1570
 Veron-Cetty, M.P., & Veron, P. 2006, *A&A*, 455, 773
 Weidinger, M., Rieger, M., Spanier, F. 2010, *ASTRA*, 6, 1

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